

Discovery of Very High Energy γ -ray emission from the extreme BL Lac object H2356-309 with H.E.S.S.

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The understanding of acceleration mechanisms in active galactic nuclei (AGN) jets and the measurement of the extragalactic background light (EBL) density are closely linked and require the detection of a large sample of VHE emitting extragalactic objects at varying redshifts. We report here on the discovery with the H.E.S.S. (High Energy Stereoscopic System) Atmospheric-Cherenkov telescopes of the very high energy (VHE) γ -ray emission from H 2356–309, an extreme BL Lac object located at a redshift of 0.165. The observations of this object, which was previously proposed as a southern-hemisphere VHE candidate source, were performed between June and December 2004. The total exposure is 39 hours live time, after data quality selection, which yields the detection of a signal at the level of 11.9σ (standard deviations).

1. Introduction

Blazars are active galactic nuclei possessing relativistic jets pointing toward the observer. Their broad-band spectral energy distributions (SED), characterised by two large peaks, are dominated by non-thermal emission, which is believed to be produced in these jets. The low energy peak is commonly associated with synchrotron emission from relativistic electrons in local magnetic fields, while the high energy peak may have a leptonic or hadronic origin. In leptonic models, the γ -rays are produced in inverse Compton interactions between these electrons and local or external photons. Whereas in hadronic models they are produced in π_0 decays resulting from the interaction of relativistic protons with the ambient radiation fields or matter.

Among blazars, BL Lacertae objects are characterised by the absence of emission line features in the optical part of their SED. In the past, according to results from initial radio and X-ray surveys, these objects were separated in two families, the low-frequency-peaked BL Lac objects (LBL) or the high-frequency-peaked BL Lac objects (HBL), depending on whether their synchrotron emission is stronger in the radio-IR band or in the UV-X band. Recently, different surveys have shown that the synchrotron peak may occur at any energy between radio and X-ray domain. The currently accepted unified picture of blazars [1] suggests that those objects with lower luminosities and higher frequency peaked synchrotron emission are efficient particle accelerators at very high energies and emit VHE γ -rays.

The detection and characterisation of the spectral properties of such objects is crucial for the understanding of the underlying acceleration mechanisms. It is also a promising tool for an indirect measurement of the diffuse EBL density, which carries unique cosmological information about galaxy formation and evolution, but whose direct measurement is subject to large uncertainties. The EBL photons in the range 0.1–20 micron μm have the appropriate energy to pair-produce with sub-TeV to TeV γ -rays, which are thus absorbed. Hence the signature of this absorption on the VHE spectra of blazars, which is a function of the redshift of the source, can be used as a probe of EBL density, if the intrinsic spectra are known. Thus, the understanding of the acceleration

mechanisms and the measurement of the EBL density are closely linked and require the detection of a large sample of VHE emitting BL-Lac objects at varying redshifts. Until the present, only a small number of such sources have been discovered (see table in [2]). We report in this paper on the discovery, with the H.E.S.S. telescopes, of a new VHE γ -ray emitter, the extreme BL-Lac object H 2356–309.

H 2356–309 was first reported as the optical counter part [3] of the X-ray source 4U0009-33, discovered by the satellite UHURU in the early 1970' [4]. It was later associated with the X-ray source 1H2351-315 detected by the Large Area Sky Survey Experiment on board the HEAO satellite [5]. The host galaxy has been resolved as elliptical by different experiments [6][7][8][9]. The detection of stellar features by Falomo et al. [6] allowed the determination of its redshift, $z=0.165$. VLA and VLBA measurements have shown weak radio-emission [10][11]. The “extreme blazars” selected on the basis of their radio, optical and X-ray fluxes by Ghisellini [12], included H 2356–309 as a potential hard X-ray and TeV emitter. Subsequent observations carried out with BeppoSAX revealed a synchrotron peak at 1.8 keV [13], confirming this object as a strong candidate for TeV emission [14].

2. Observations with the H.E.S.S telescopes

The H.E.S.S. (High Energy Stereoscopic System) experiment is an array of four imaging atmospheric Cherenkov telescopes located in Namibia (23.3° S, 16.5° E, 1835 m a.s.l.). It is dedicated to the search for γ -ray emission above 100 GeV. Each telescope is equipped with a 13 m diameter dish (focal length 15 m, mirror area 107 m^2) and a 960-pixel fast camera that providing a 5° field of view (f.o.v.). See [15] for more details about H.E.S.S..

The observations of H 2356–309 were performed between June and December 2004 with the full four-telescope array. Only the events for which at least two telescopes had triggered were selected by the hardware central

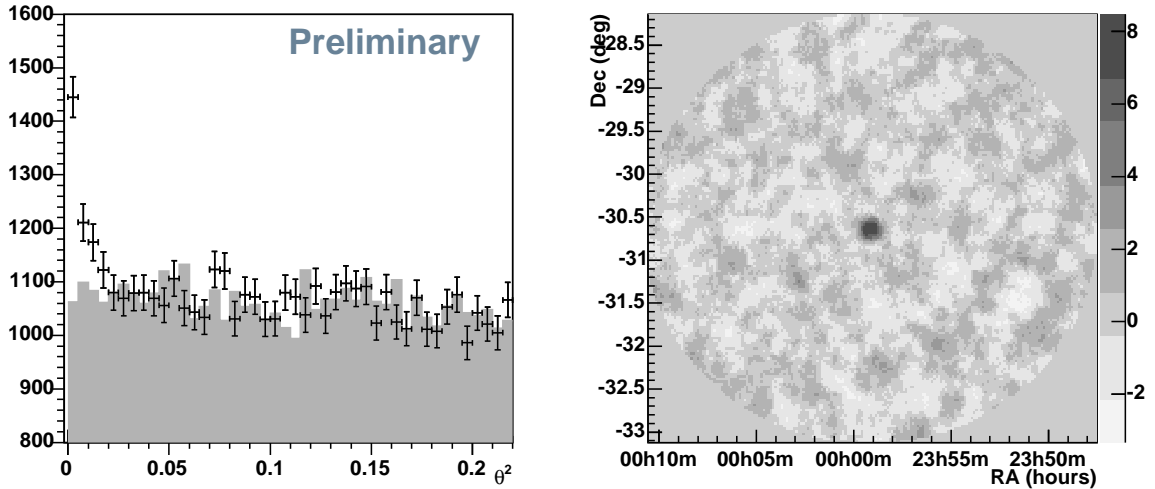


Figure 1. Left: The on-source (points) and normalized off-source background (grey filled) distributions of θ^2 . The excess at $\theta^2 < 0.02$ indicates VHE emission from the direction of H 2356–309. Right: A correlated two-dimensional significance map. The shape of the uncorrelated excess is consistent with that of a point-like emitter located at the position of H 2356–309. (preliminary).

trigger system [16]. The source has been tracked at zenith angles ranging from 7° to 54° with an average value at 20° . The data were taken in 28 minute runs using “Wobble” mode, where the source direction was positioned at $\pm 0.5^\circ$ relative to the centre of the f.o.v. of view during observations. This increases the duty-cycle of the telescope by allowing on-source observations and simultaneous estimation of the background induced by charged cosmic rays using other parts of the f.o.v.. The total exposure time is 38.9 hours after a run quality selection that removes runs for which the weather conditions are poor or the telescopes are not operating within specified requirements (trigger rate stability, hardware status).

3. Results

The results presented below are based on the H.E.S.S. standard analysis method [15]. The left side of Figure 1 shows the on-source (points) and normalized off-source background (grey filled) distributions of θ^2 , where θ is the angular distance between the direction of the candidate γ -rays and the nominal position of the source. The excess at $\theta^2 < 0.02$ indicates VHE emission from the direction of H 2356–309. An excess of 845 events with $\theta^2 < 0.02 \text{ deg}^2$ is observed at a significance level of 11.9σ (likelihood method of Li & Ma [17]). A correlated two-dimensional significance map is shown on the right side of Figure 1. As expected, the shape of the uncorrelated excess is consistent with that of a point-like emitter located at the position of H 2356–309.

The observed integral flux above 200 GeV versus modified Julian date for each dark period is shown in Figures 2a to 2d. Figures 2a, 2b and 2c, respectively, show the night-by-night light curves for June-July, September-October and November-December 2004 periods. The month-by-month light curve is shown in Figure 2d. A fit of a constant to the data shown in Figure 2d yields a 0.3% χ^2 probability, indicating flux variability on monthly time scales. No evidence for significant variability is found on shorter time scales. Table 1 shows the

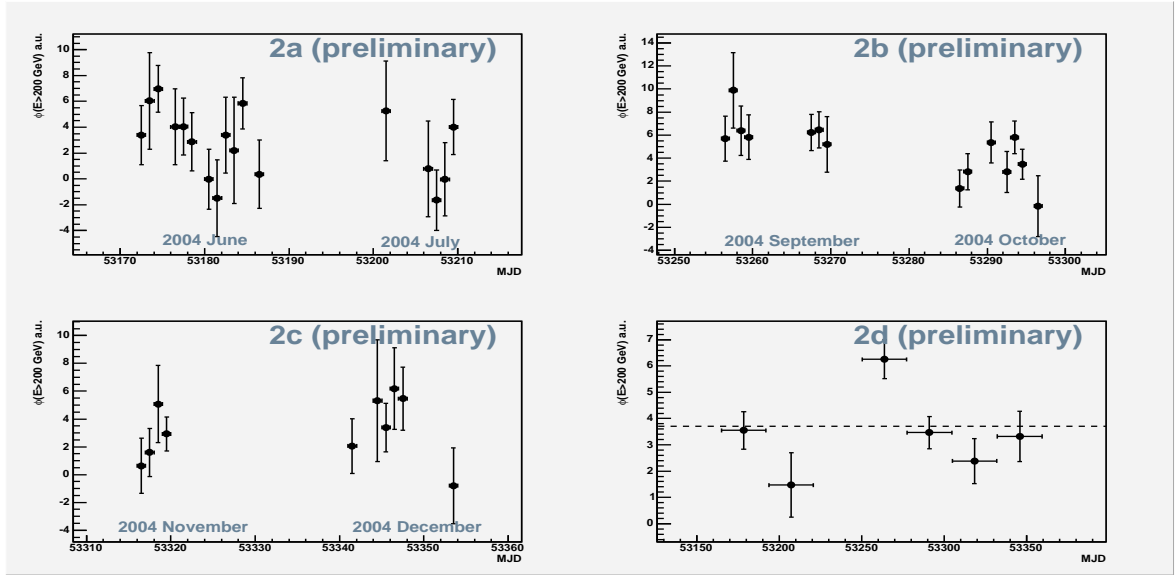


Figure 2. The panels show the night-by-night light for June-July (2a), September-October (2b) and November-December (2c) 2004 periods. The month-by-month light curve is given in the final panel, where the dashed line shows the result of a constant fit (preliminary).

Table 1. For each dark period : the number of runs, the live time, the average zenith angle, the number of *on* and *off* events passing selection cuts, the *off*-source normalisation coefficient, the excess and the corresponding significance.

Dark periods	Nruns	Obs. Time [hrs]	Zen[deg]	On	Off	Norm	Excess	Sig. [σ]
June	21	7.8	19	1156	6911	0.1418	176	5.1
July	7	3.0	10	351	2297	0.1426	23	1.2
September	16	6.8	13	1001	5166	0.1425	265	8.6
October	23	9.8	15	1241	7257	0.1422	209	5.9
November	11	3.9	20	392	2320	0.1422	62	3.1
December	19	7.6	36	810	4917	0.1423	110	3.8
Total	97	38.9	20	4951	28868	0.1422	845	11.9

results obtained during different dark periods. The source shows a maximal activity in September, for which the excess of 265 γ -rays has a significance of 8.6σ . The VHE spectrum of this source will be presented at the conference.

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